



Supporting Online Material for

The Occurrence and Mass Distribution of Close-in Super-Earths, Neptunes, and Jupiters

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Published 29 October 2010, *Science* **330**, 653 (2010)
DOI: 10.1126/science.1194854

This PDF file includes:

Materials and Methods
Figs. S1 to S7
Tables S1 to S3
References

Materials and Methods

In this supplement to “The Occurrence and Mass Distribution of Close-in Super-Earths, Neptunes, and Jupiters” by Howard et al., we provide materials and methods for our analysis. We describe the radial velocity (RV) measurements at Keck observatory (Section S1) and the Eta-Earth Survey stars and planets (Section S2). We discuss calculations to compute limits on the mass of planetary companions for each star and how these limits are folded together to compute a search completeness function, C (Section S3). Our methods for computing occurrence rates as a function of planet mass are described in Section S4. In Section S5, we calculate the number of stars detectable by the Kepler mission with orbital period $P < 50$ d and mass $1\text{--}8 M_{\text{Earth}}$.

S1 Observations and Doppler Analysis

We observed the Eta-Earth Survey stars in Table S1 using the HIRES echelle spectrometer (*I*) on the 10-m Keck I telescope. Exposure times ranged from a few seconds for the brightest stars to 500 seconds for the faintest. All observations were made with an iodine cell mounted directly in front of the spectrometer entrance slit. The dense set of molecular absorption lines imprinted on the stellar spectra provide a robust wavelength fiducial against which Doppler shifts are measured, as well as strong constraints on the shape of the spectrometer instrumental profile at the time of each observation (2, 3).

We measured the Doppler shift from each star-times-iodine spectrum using a forward modeling procedure modified from the method described in (4). The most significant modification is the way we model the intrinsic stellar spectrum, which serves as a reference point for the relative Doppler shift measurements for each observation. Butler et al. use a version of the Jansson deconvolution algorithm (5) to remove the spectrometer’s instrumental profile from an iodine-free template spectrum. We instead use a deconvolution algorithm that employs a more effective regularization scheme, which results in significantly less noise amplification and improved Doppler precision.

All of the measurements for this analysis were made after 2004 Aug. when HIRES was upgraded with a new CCD and started to consistently achieve 1 m s^{-1} precision (6). Many of the Eta-Earth Survey stars also have “pre-upgrade” measurements of lower quality that are not used here.

For each observation we also measure S_{HK} and the related quantity $\log R'_{\text{HK}}$ from the flux in the cores of the Ca II H and K lines in the HIRES spectra (7). These indicators of chromospheric activity provide useful checks that detected RV variations are due to a planet in orbital motion and not magnetic activity on the stellar surface modulated by stellar rotation and activity cycles. We identified 13 stars with significant correlations between RV and S_{HK} . For these stars, we de-correlated the activity signal by subtracting from the RVs a linear fit to the measured RVs as

a function of S_{HK} . Typically this process removes $2\text{--}4\text{ m s}^{-1}$ of RV variability.

S2 Eta-Earth Survey Target Stars and Planets

We selected 235 G, K, and M-type dwarf stars for the NASA-UC Eta-Earth Survey. This survey was designed as a search for planets orbiting a well-controlled sample of stars with the goal of characterizing the population of planets using detections and non-detections. The stars were selected from the Hipparcos catalog (8) based on brightness ($V < 11$), distance ($< 25\text{ pc}$), luminosity ($M_V > 3.0$), chromospheric activity ($\log R'_{\text{HK}} < -4.7$), lack of stellar companions, and observability from Keck Observatory. The resulting set of stars is nearly free of selection bias; in particular, stars were neither included nor excluded based on their likelihood to harbor a planet. Metallicity was not used as a selection criterion. Here we focus on the 166 G and K-type stars with masses in the range $0.54\text{--}1.28$ solar masses and $B - V < 1.4$. Table S1 lists the names of these stars (using HD or Hipparcos identifiers), their spectral types, masses, and the number of observations for each star. Figure S1 shows the distribution of stellar masses.

Each star was observed over 5 years and during at least one “high-cadence run” of 6–12 observations in a 12 night span to increase our sensitivity to short-period signals. For some stars with candidate planets, we made 2–5 consecutive observations and averaged the resulting velocities for a single measurement. Figure S2 shows a histogram of the number of RV measurements per star. We made at least 20 RV (and as many as 144) measurements for each star; the median number of measurements is 33. Stars with substantially more than the median number of observations typically have detected planets or particularly low astrophysical jitter. Additionally, when a candidate planet with a formal false alarm probability (FAP) (6) of $< 20\%$ was identified, the host star was observed intensively as telescope time permitted.

Table S2 lists the 33 planets detected around 22 Eta-Earth Survey stars, their orbital periods, minimum masses, and references in the literature. We use HD names to identify the stars hosting each planet to match the star names in Table S1. Some of these planets were initially detected by other groups, but were confirmed by our HIRES measurements. We do not announce any new planets here. Table S3 lists the orbital periods and minimum masses ($M \sin i$) of the 12 candidate planets orbiting Eta-Earth Survey stars. The candidates have formal FAPs of $< 5\%$. We typically publish planets when their FAPs drop below 1% and we are sure that the periodic signal is due to a planet and not astrophysical noise or systematic errors.

The 33 detected planets and 12 candidate planets (45 planets in total) orbit 32 unique stars. Eight of the 32 stars have multiple detected planets (25% multiplicity). Of the 16 stars with a detected planet of any mass in a $P < 50\text{ d}$ orbit, seven have multiple detected planets (44% multiplicity).

S3 Search Completeness

Every star in the Eta-Earth Survey has a different observing history. Figure S3 shows the RV time series for a star with a large number of RV measurements (HD 185144, top) and for a star with a more sparse set of observations (HD 84737, bottom).

Figure S4 shows the distribution of velocity RMS for Eta-Earth Survey stars, excluding stars with detected planets (Table S2). Long-term RV trends were fit for and subtracted before computing the velocity RMS. Activity correlations were also subtracted for 13 stars before computing the RMS.

For each star we estimate the maximum $M \sin i$ value compatible with the RV measurements as a function of prospective orbital period. On a fine grid of orbital periods (2000 log-spaced trial orbital periods per \log_{10} period interval) we fit the RVs to a circular orbit using a partially-linearized, least-squares fitting procedure (9) that minimized χ^2 . For each fit, the orbital period of the trial planet was allowed to vary within the narrow period bin. For stars with long-term RV trends, we simultaneously fit for these objects with each trial planet.

Each fit produced a best-fit Doppler semi-amplitude K at the prospective orbital period P . We transformed $K(P)$ into $M \sin i_{\max}(P) = M_m(P)$ using masses for each star (10, 11). The black lines in Figure S5 trace $M_m(P)$ for two Eta-Earth Survey stars. For a well-observed star like HD 185144, we are sensitive to planets with masses as low as $M \sin i \approx 1 M_{\text{Earth}}$ for $P < 3$ d. For a more typical star with fewer observations (e.g., HD 84737), our measurements have sufficient precision to detect nearly all planets with $M \sin i > 10 M_{\text{Earth}}$ and $P < 50$ d.

For short periods, $M_m(P)$ is a rapidly varying function of P because our observations span several years, rendering nearby short periods distinguishable. This high period resolution is useful for testing for the existence of planets with specific orbital periods, but for our purpose these fine details are not helpful. So we compute $M_{m,e}(P)$, the upper envelope of $M_m(P)$ on a coarser period grid with 20 log-spaced trial orbital periods per \log_{10} period interval. $M_{m,e}(P)$ is the maximum value of $M_m(P)$ within each coarse period interval. The green lines in Figure S5 show that $M_{m,e}(P)$ captures the behavior of $M_m(P)$ on a scale that is more appropriate for computing a search completeness correction.

Using $M_{m,e}(P)$ for the set of stars without detected or candidate planets we compute the search completeness function, $C(P, M \sin i)$, depicted in Figure 1 of the main paper. $C(P, M \sin i)$ measures the fraction of stars with sufficient measurements to rule out a planet of a given minimum mass and orbital period. We compute this quantity on the coarse period grid used for $M_{m,e}(P)$ and on a coarse grid of $M \sin i$ with 20 log-spaced $M \sin i$ bins per \log_{10} interval. At a given P and $M \sin i$, $C(P, M \sin i)$ is equal to the fraction of stars with $M_{m,e}(P) < M \sin i$. That is, $C(P, M \sin i)$ measures the fraction of stars where the limit on $M \sin i$ is below a certain value, for a given orbital period.

S4 Planet Occurrence

To compute planet occurrence rates, we counted the number of stars hosting detected and candidate planets in each of five mass domains with $P < 50$ d: $M \sin i = 3\text{--}10$, $10\text{--}30$, $30\text{--}100$, $100\text{--}300$, and $300\text{--}1000 M_{\text{Earth}}$. Multiple planets in the same mass domain that orbit the same star were only counted once, for the host star (e.g., HD 69830 is only counted once for the $M \sin i = 10\text{--}30 M_{\text{Earth}}$ domain despite hosting two planets in that mass range). Planet host stars can be counted in multiple mass domains though (e.g., 55 Cnc hosts planets in the $M \sin i = 3\text{--}10$, $3\text{--}100$, and $100\text{--}300 M_{\text{Earth}}$ domains and is counted in each of them). In addition to the five mass domains above, we also computed the planet occurrence rate in a domain with $M \sin i = 3\text{--}30 M_{\text{Earth}}$ and $P < 50$ d for comparison with the estimates by Mayor et al.

Our estimates of planet occurrence depend on a “missing planet correction” computed for each detected and candidate planet. This correction is a statistical estimate of the number of planets of similar P and $M \sin i$ that remain undetected in our sample. To compute this effective number of planets for each detected and candidate planet, we add $n_{\text{miss}} = 1/C(P, M \sin i) - 1$ missed planets. Since $C = 1$ over the three highest mass domains (see Figure 1), $n_{\text{miss}} = 0$ for these domains. For the lower right corner of the $M \sin i = 10\text{--}30 M_{\text{Earth}}$ domain, we have $C < 1$ and compute $\Sigma n_{\text{miss}} = \Sigma(1/C(P, M \sin i) - 1) = 4.6$ missed planets by summing over detected and candidate planets in this domain. For much of the $M \sin i = 3\text{--}10 M_{\text{Earth}}$ domain $C < 1$ and we compute $\Sigma n_{\text{miss}} = 10.2$ missed planets with a similar summation.

The best-fit values and error bars on planet occurrence and related parameters are computed using binomial statistics. We compute the probability distribution functions (pdfs) of the occurrence rate in each bin as the probability of drawing $n_{\text{det}} + n_{\text{cand}}$ detected and candidate planets out of n_{star} stars. We multiply these occurrence rates and associated uncertainties by $(n_{\text{det}} + n_{\text{cand}} + n_{\text{miss}})/(n_{\text{det}} + n_{\text{cand}})$ to account for n_{miss} missed planets. The median and 68.3% confidence intervals of these distributions are reported as the best-fit occurrence values and their “1- σ ” errors.

We computed the best-fit power-law model to planet occurrence as a function of planet mass by randomly drawing occurrence rates from the pdfs of the five mass domains. For each of 30,000 trials we fit the occurrence rates to a power law model of the form $df/d \log M = kM^\alpha$. The resulting distribution of k and α is shown in Figure S6. Each trial specifies a particular power law and predicts the occurrence rate of planets in several important mass domains. Figure S7 shows the occurrence rate pdfs in two important mass domains, $0.5\text{--}2.0 M_{\text{Earth}}$ (“Earth-like” planets) and $1\text{--}8 M_{\text{Earth}}$ (“Kepler planets”; see below).

Our analysis makes several assumptions worth noting. First, we restrict orbital fits to circular when computing $M_m(P)$. This reduces computational complexity and is physically well-motivated for close-in planets tidally coupled to their host stars. The detected and candidate close-in planets in the $M \sin i = 3\text{--}10 M_{\text{Earth}}$ domain all have orbits that are consistent with circular or have $e < 0.2$. RV signals from slightly eccentric orbits are well approximated by circular orbits, and thus any reductions in sensitivity as measured by $M_{m,e}(P)$ and $C(P, M \sin i)$ will be small for such planets. We also note that our limits on $M \sin i$ are for a single planet. For

stars with multiple undetected planets with nearly identical K values, our sensitivity is reduced from the limits expressed by $M_{m,e}(P)$. Such systems are likely rare though since this scenario requires a compensation of P and $M \sin i$ to produce nearly identical K . We searched for such systems in our existing detections and found that only three of eleven stars with a candidate or detected planet with $M \sin i = 3\text{--}30 M_{\text{Earth}}$ and $P < 50$ d had a second planet in the same mass and period range with a Doppler semi-amplitude within a factor of $\sqrt{2}$ of the first planet. A further approximation is the binary nature of $M_m(P)$ and $M_{m,e}(P)$; planets are deemed detectable above this limit and undetectable below it. This sharp dividing line is averaged out though by using all 166 Eta-Earth stars to compute C . Our planet occurrence measurements are very weakly dependent on choices in constructing the fine and coarse period grids of $M_m(P)$ and $M_{m,e}(P)$. Finally, errors in n_{miss} for a given detected or candidate planet scale as $1/C$ and can be large where the search is relatively incomplete and C is small.

We considered the impact of a 10% error in C , resulting from one of the above approximations. For this test we used a modified search completeness $C'(P, M \sin i) = 1.1 \times C(P, M \sin i)$ and ran the statistical analysis again. (The value of C' was capped at 1.0.) With this change, the number of missed planets, n_{miss} , in the $M \sin i = 3\text{--}10$ and $10\text{--}30 M_{\text{Earth}}$ bins goes from 10.2 and 4.6 planets to 8.5 and 3.6 planets, respectively. Thus a 10% change in C results in 17% and 22% changes in n_{miss} for the two bins. However, this effect is diluted by the detected and candidate planets in each bin, which are held fixed in this numerical experiment. The resulting effect on the planet occurrence rates is small. For $M \sin i = 3\text{--}10 M_{\text{Earth}}$, the planet occurrence rate goes from $11.8^{+4.3}_{-3.5}\%$ to $10.7^{+3.9}_{-3.2}\%$ and for $M \sin i = 10\text{--}30 M_{\text{Earth}}$ the occurrence rate goes from $6.5^{+3.0}_{-2.3}\%$ to $5.8^{+2.7}_{-2.1}\%$. The power law parameter change from $k = 0.39^{+0.27}_{-0.16}$, and $\alpha = -0.48^{+0.12}_{-0.14}$ to $k = 0.34^{+0.23}_{-0.14}$, and $\alpha = -0.46^{+0.13}_{-0.15}$. Thus, a 10% in the search completeness C results in changes in the measured planet occurrence rates that are small compared to their statistical errors.

S5 Kepler

It would be desirable to compare our planet occurrence measurements in the planet mass/orbital period plane with occurrence measurements by the Kepler mission (12) in the planet radius/orbital period plane. The current Kepler data release supports the qualitative observation that close-in small planets are much more common than large ones (13). Unfortunately, meaningful quantitative comparisons are not yet possible because the data are currently sequestered for 400 of the 706 stars identified as planet hosts. For the planet candidates with released data, only planet radii are available, not masses. Moreover, the initial data release only covers a 43.2 d span (robust transit detection typically requires three or more observed transit events), rendering highly incomplete planet occurrence statistics for orbital periods out to 50 d. We therefore proceed with a prediction for the number of small planets that Kepler will detect.

To estimate the number of planets detected by the Kepler mission with $P < 50$ d and radii of 1–2 Earth-radii (implying masses of 1–8 M_{Earth} assuming Earth density of 5.5 kg m^{-3}), we

must account for the period dependence of the planet occurrence rate and of the probability of a planet transiting its host star (rendering it potentially detectable by Kepler). Our power law model of planet occurrence predicts 22% occurrence for $P < 50$ d and $M \sin i = 1\text{--}8 M_{\text{Earth}}$. We naively assume that the distribution of these planets in orbital period is uniform in $\log P$ over $P = 1\text{--}50$ d. For a circular orbit, the probability that a planet transits depends on its orbital period and is given by $p_{\text{transit}} = R_{\text{star}}/a = (4\pi^2 R_{\text{star}}^3 / GM_{\text{star}} P^2)^{1/3}$, where M_{star} and R_{star} are the stellar mass and radius, a and P are the semimajor axis and orbital period of the planet, and G is the gravitational constant. We naively assign $R_{\text{star}} = R_{\text{sun}}$ for all stars in this calculation. Integrating the transit probability and planet occurrence over $P = 1\text{--}50$ d, we find that 1.2–2.6% of stars (68% confidence interval) will harbor a transiting planet with the specified parameters. We estimate that the $\sim 10^4$ G and K dwarfs brighter than 13th magnitude (14) are bright enough to yield a transit detection down to 1 Earth-radius. Thus we estimate that 120–260 stars host planets with radii of 1–2 Earth-radii and $P < 50$ d that are detectable by Kepler.

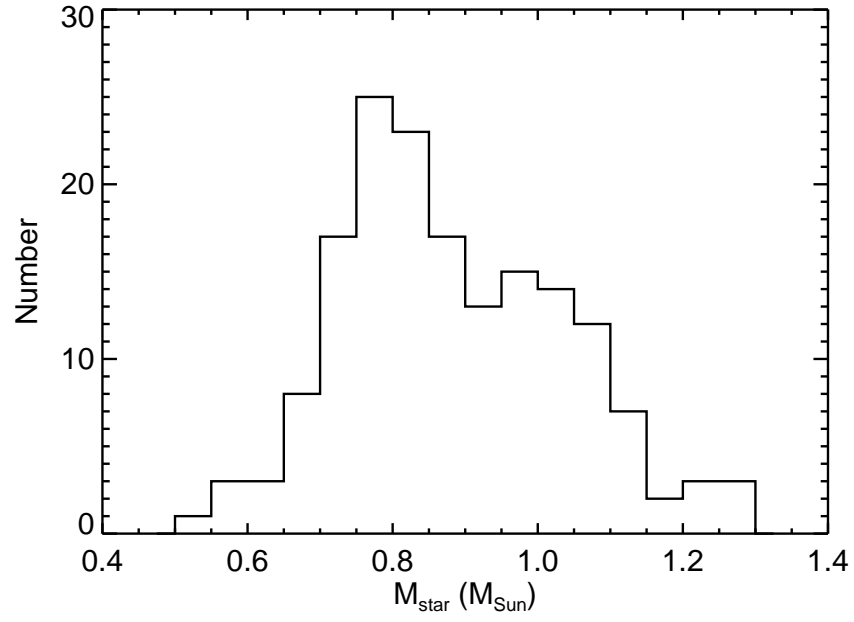


Figure S1: Histogram of stellar masses for Eta-Earth stars.

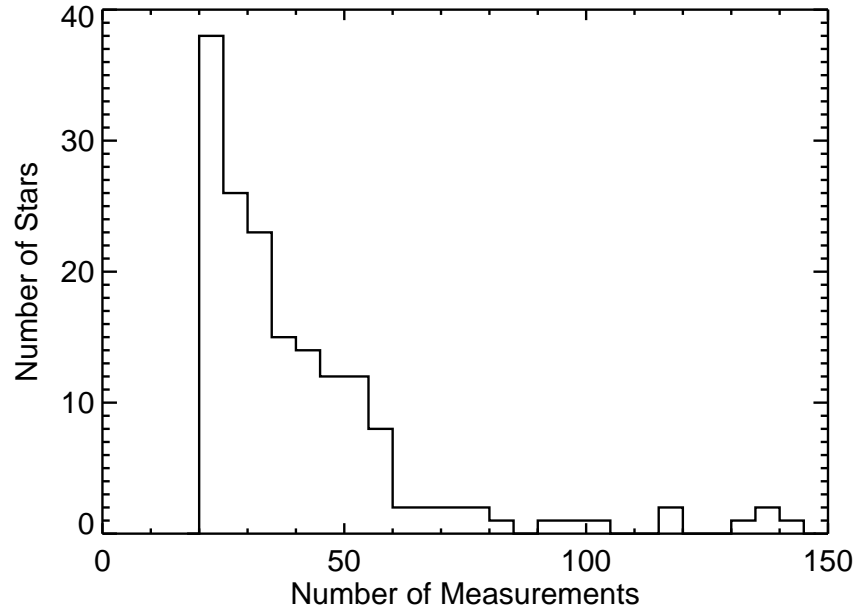


Figure S2: Histogram of number of RV measurements per Eta-Earth star.

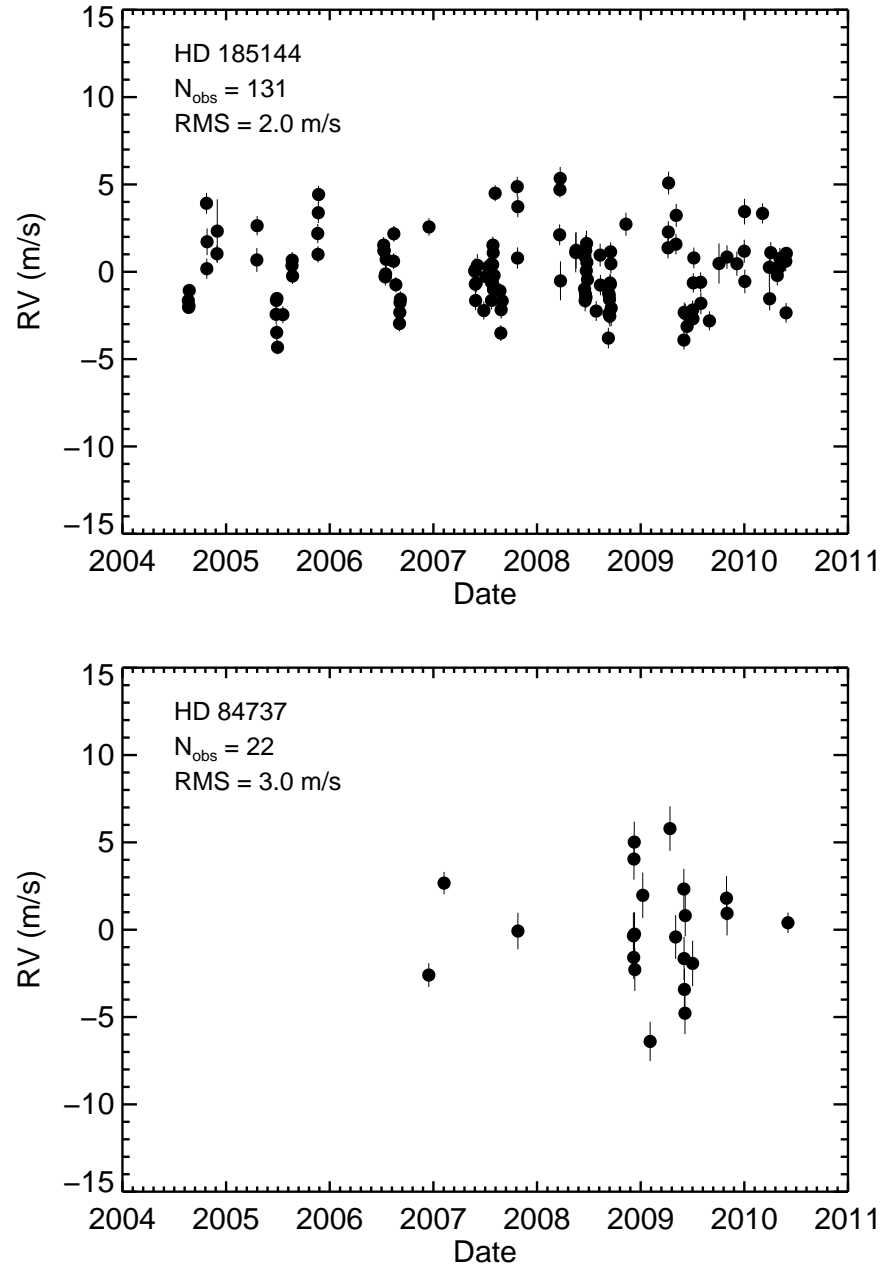


Figure S3: Radial velocity time series for a well-observed star (HD 185144, top) and a star with fewer measurements (HD 84737, bottom). The star name, number of observations, and velocity RMS are indicated in each panel.

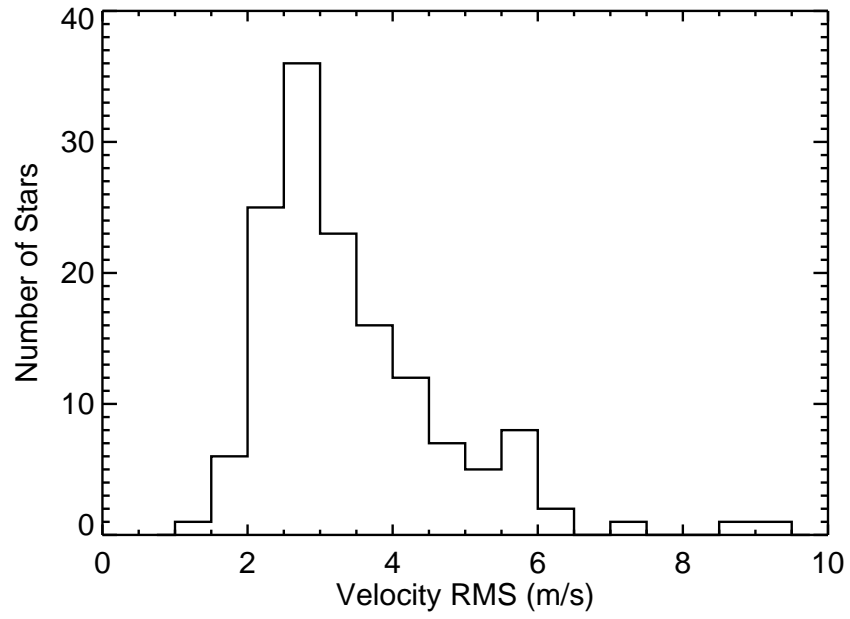


Figure S4: Histogram of the RMS of the RV measurements for each star. Stars with detected planets (Table S2) are excluded.

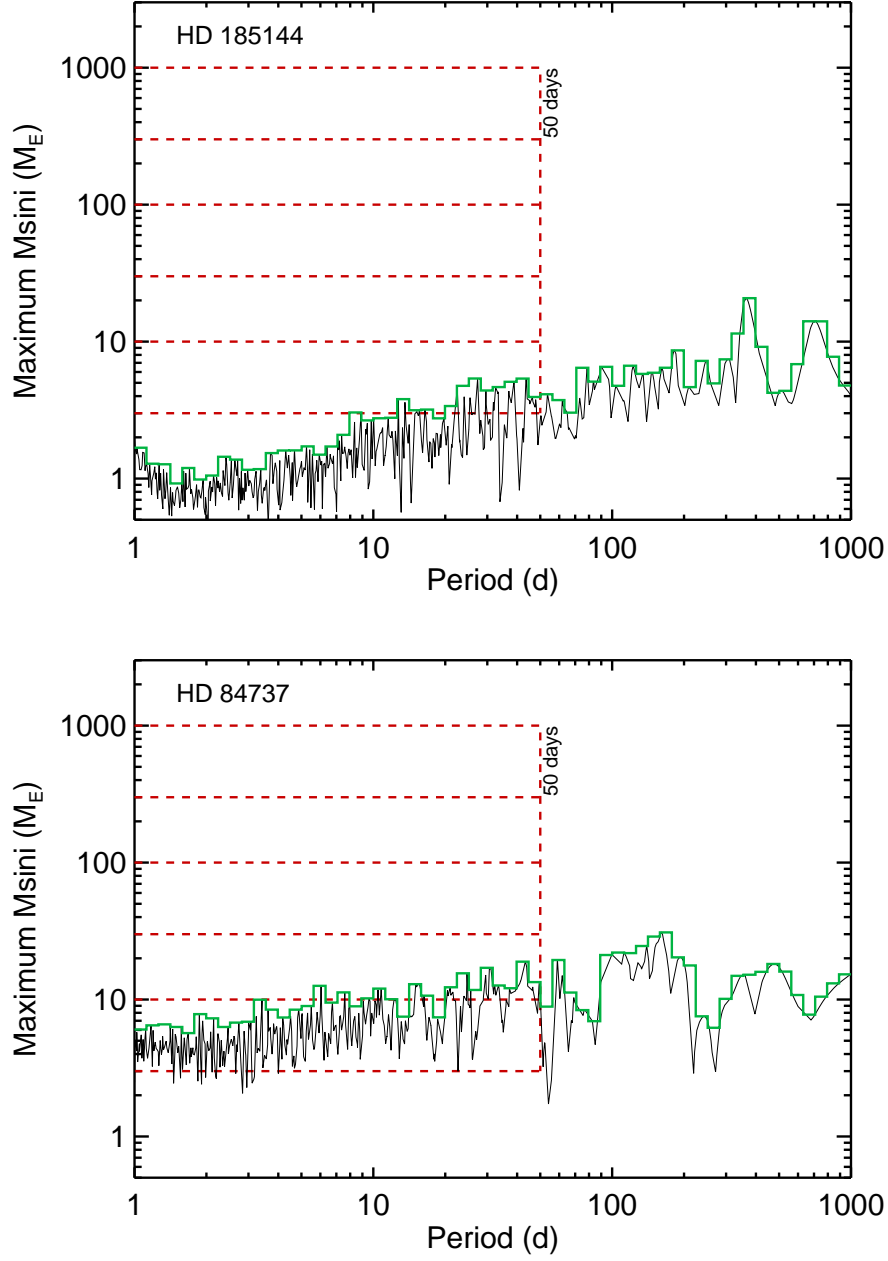


Figure S5: Limits on $M \sin i$ as a function of orbital period for the stars whose RV time series are plotted in Figure S3. The thin black lines trace this limit, $M_m(P)$, on a fine period grid. The thick green lines, $M_{m,e}(P)$, trace the upper envelope of $M_m(P)$ on a coarser period grid. Five mass domains are indicated by dashed red lines for planets with $P < 50$ d and $M \sin i = 3\text{--}10$, $10\text{--}30$, $30\text{--}100$, $100\text{--}300$, and $300\text{--}1000 M_{\text{Earth}}$.

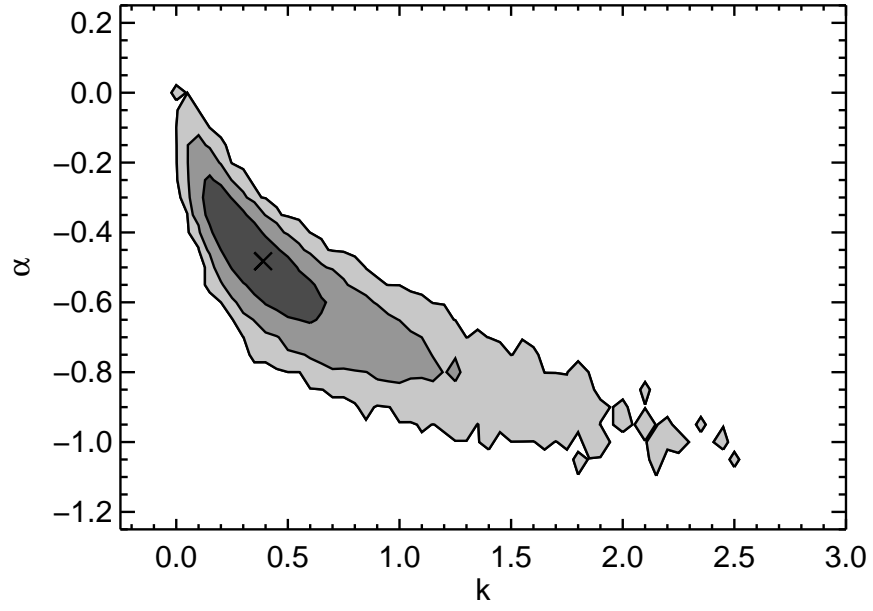


Figure S6: Probability distributions for the parameters k and α of a power law model of planet occurrence, $df/d \log M = kM^\alpha$. The X marks the best-fit values and the contours represent 68.3%, 95.4%, and 99.7% confidence intervals (1, 2, and 3- σ).

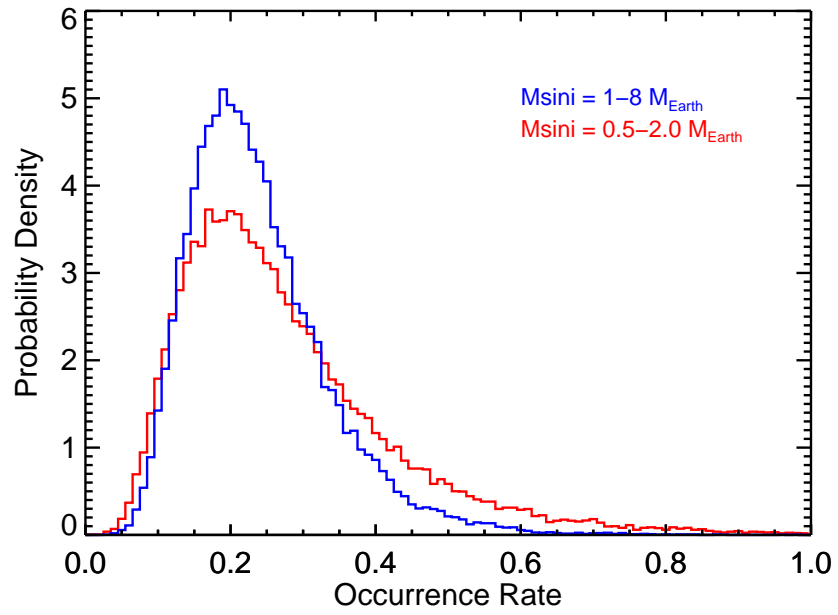


Figure S7: Probability distributions for the occurrence rate in two mass domains based on a power law model.

Table S1. G and K-type Target Stars in the Eta-Earth Survey

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 1461	G0	1.08	154
HD 3651	K0	0.89	29
HD 3765	K2	0.84	35
HD 4256	K2	0.85	36
HD 4614	G0	0.99	30
HD 4614 B	K7	0.57	28
HD 4628	K2	0.72	49
HD 4747	G8	0.82	22
HD 4915	G0	0.90	37
HD 7924	K0	0.83	135
HD 9407	G6	0.98	97
HD 10476	K1	0.83	56
HD 10700	G8	0.95	133
HD 12051	G5	0.99	52
HD 12846	G2	0.88	36
HD 14412	G5	0.78	37
HD 16160	K3	0.76	47
HD 17230	K5	0.69	31
HD 18143	G5	0.90	35
HD 18803	G8	1.00	32
HD 19373	G0	1.20	47
HD 20165	K1	0.82	26
HD 20619	G1	0.91	35
HD 22879	F9	0.79	22
HD 23356	K2	0.78	22
HD 23439	K1	0.67	26
HD 24238	K0	0.73	29
HD 24496	G0	0.94	47
HD 25329	K1	0.83	34
HD 25665	G5	0.78	21
HD 26965	K1	0.78	41

Table S1—Continued

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 29883	K5	0.76	23
HD 32147	K3	0.83	52
HD 32923	G4	1.03	26
HD 34721	G0	1.12	21
HD 34411	G0	1.13	40
HD 36003	K5	0.73	42
HD 37008	K2	0.73	22
HD 38230	K0	0.83	24
HD 38858	G4	0.92	35
HD 40397	G0	0.92	23
HD 42618	G4	0.96	59
HD 45184	G2	1.04	46
HD 48682	G0	1.17	27
HD 50692	G0	1.00	37
HD 51419	G5	0.86	40
HD 51866	K3	0.78	32
HD 52711	G4	1.02	46
HD 55575	G0	1.26	32
HD 62613	G8	0.94	24
HD 65277	K5	0.72	21
HD 65583	G8	0.76	26
HD 68017	G4	0.85	43
HD 69830	K0	0.87	46
HD 72673	K0	0.78	23
HD 73667	K1	0.72	22
HD 75732	G8	0.91	96
HD 84035	K5	0.73	22
HD 84117	G0	1.15	22
HD 84737	G0	1.22	24
HD 86728	G3	1.08	28
HD 87883	K0	0.80	30

Table S1—Continued

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 89269	G5	0.89	29
HD 90156	G5	0.90	28
HD 92719	G2	1.10	24
HD 95128	G1	1.08	22
HD 97101	K8	0.60	21
HD 97343	G8	0.89	35
HD 97658	K1	0.78	61
HD 98281	G8	0.85	46
HD 99491	K0	1.01	71
HD 99492	K2	0.86	47
HD 100180	G0	1.10	24
HD 100623	K0	0.77	32
HD 103932	K5	0.76	44
HD 104304	G9	1.02	23
HD 109358	G0	1.00	41
HD 110315	K2	0.70	37
HD 110897	G0	1.23	29
HD 114613	G3	1.28	21
HD 114783	K0	0.86	45
HD 115617	G5	0.95	61
HD 116442	G5	0.76	25
HD 116443	G5	0.73	55
HD 117176	G4	1.11	30
HD 120467	K4	0.71	20
HD 122064	K3	0.80	43
HD 122120	K5	0.71	36
HD 125455	K1	0.79	20
HD 126053	G1	0.86	30
HD 127334	G5	1.10	24
HD 130992	K3	0.77	36
HD 132142	K1	0.77	21

Table S1—Continued

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 136713	K2	0.84	79
HD 139323	K3	0.89	91
HD 140538 A	G2	1.06	58
HD 141004	G0	1.14	68
HD 143761	G0	1.00	29
HD 144579	G8	0.75	30
HD 145675	K0	1.00	59
HD 145958 A	G8	0.91	44
HD 145958 B	K0	0.88	31
HD 146233	G2	1.02	52
HD 146362 B	G1	1.07	29
HD 148467	K5	0.67	22
HD 149806	K0	0.94	28
HD 151288	K5	0.59	22
HD 151541	K1	0.83	29
HD 154088	G8	0.97	67
HD 154345	G8	0.88	53
HD 154363	K5	0.64	25
HD 155712	K0	0.79	39
HD 156668	K2	0.77	93
HD 156985	K2	0.77	34
HD 157214	G0	0.91	25
HD 157347	G5	0.99	46
HD 158633	K0	0.78	20
HD 159062	G5	0.94	29
HD 159222	G5	1.04	55
HD 161797	G5	1.15	22
HD 164922	K0	0.94	50
HD 166620	K2	0.76	35
HD 168009	G2	1.02	24
HD 170493	K3	0.81	33

Table S1—Continued

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 172051	G5	0.87	28
HD 176377	G0	0.92	32
HD 179957	G4	1.01	39
HD 179958	G4	1.03	38
HD 182488	G8	0.96	45
HD 182572	G8	1.14	27
HD 185144	K0	0.80	122
HD 185414	G0	1.07	27
HD 186408	G1	1.07	35
HD 186427	G3	0.99	44
HD 190067	G7	0.80	50
HD 190360	G6	1.01	45
HD 190404	K1	0.70	21
HD 190406	G1	1.09	32
HD 191785	K1	0.83	22
HD 191408	K3	0.69	36
HD 192310	K0	0.82	45
HD 193202	K5	0.67	38
HD 196761	G8	0.83	27
HD 197076	G5	0.99	86
HD 201091	K5	0.66	64
HD 201092	K7	0.54	62
HD 202751	K2	0.75	42
HD 204587	K5	0.68	20
HD 208313	K0	0.80	23
HD 210277	G0	1.01	49
HD 210302	F6	1.28	23
HD 213042	K5	0.74	37
HD 215152	K0	0.78	27
HD 216520	K2	0.83	60
HD 216259	K0	0.69	50

Table S1—Continued

Name	Spec. Type	Mass (M_{\odot})	Num. Obs.
HD 217014	G2	1.09	26
HD 217107	G8	1.10	41
HD 218868	K0	0.99	53
HD 219134	K3	0.78	74
HD 219538	K2	0.81	30
HD 219834 B	K2	0.82	24
HD 220339	K2	0.73	36
HD 221354	K2	0.85	79
HIP 18280	K7	0.59	22
HIP 19165	K4	0.70	21
HIP 41689	K7	0.62	20

Table S2. Detected Planets in the Eta-Earth Survey

Planet	Star	Period (d)	$M \sin i$ (M_{\oplus})	Reference
14 Her b	HD 145675	1754	1651	(15)
16 Cyg b	HD 186427	798	521	(15)
47 UMa b	HD 95128	1090	826	(16)
47 UMa c	HD 95128	2590	252	(16)
51 Peg b	HD 217014	4.2	147	(17)
55 Cnc b	HD 75732	14.7	264	(18)
55 Cnc c	HD 75732	44.4	53.4	(18)
55 Cnc d	HD 75732	5371	1241	(18)
55 Cnc e	HD 75732	2.8	7.6	(18)
55 Cnc f	HD 75732	261	46.3	(18)
61 Vir b	HD 115617	4.2	5.1	(19)
61 Vir c	HD 115617	38.0	11	(19)
61 Vir d	HD 115617	123	23	(19)
70 Vir b	HD 117176	116	2372	(20)
HD 1461 b	HD 1461	5.8	8	(21)
HD 3651 b	HD 3651	62.2	72.8	(22)
HD 7924 b	HD 7924	5.5	9.3	(6)
HD 69830 b	HD 69830	8.7	10.2	(23)
HD 69830 c	HD 69830	31.6	11.9	(23)
HD 69830 d	HD 69830	197	17.9	(23)
HD 87883 b	HD 87883	2754	558	(24)
HD 90156 b	HD 90156	49.6	16.7	(25)
HD 99492 b	HD 99492	17.0	33.7	(15)
HD 114783 b	HD 114783	493	351	(26)
HD 154345 b	HD 154345	3341	304	(27)
HD 156668 b	HD 156668	4.6	4.1	(28)
HD 164922 b	HD 164922	1155	114	(15)
HD 190360 b	HD 190360	2915	497	(29)
HD 190360 c	HD 190360	17.1	18.7	(29)
HD 210277 b	HD 210277	442	405	(15)
HD 217107 b	HD 217107	7.1	443	(30)

Table S2—Continued

Planet	Star	Period (d)	$M \sin i$ (M_{\oplus})	Reference
HD 217107 c	HD 217107	4270	831	(30)
ρ CrB b	HD 143761	39.8	338	(15)

Table S3. Candidate Planets in the Eta-Earth Survey

Name	Period (d)	$M \sin i$ (M_{\oplus})
Candidate 1	18.1	6.5
Candidate 2	15.3	7.3
Candidate 3	9.5	8.1
Candidate 4	13.7	8.9
Candidate 5	15.0	12
Candidate 6	67	15
Candidate 7	74	20
Candidate 8	71	33
Candidate 9	777	35
Candidate 10	338	46
Candidate 11	1219	62
Candidate 12	1697	94

References

1. S. S. Vogt, *et al.*, *Proc. SPIE Instrumentation in Astronomy VIII*, David L. Crawford; Eric R. Craine; Eds., Vol. 2198, p. 362 (1994).
2. G. W. Marcy, R. P. Butler, *Proceedings of the Astronomical Society of the Pacific* **104**, 270 (1992).
3. J. A. Valenti, R. P. Butler, G. W. Marcy, *Proceedings of the Astronomical Society of the Pacific* **107**, 966 (1995).
4. R. P. Butler, *et al.*, *Proceedings of the Astronomical Society of the Pacific* **108**, 500 (1996).
5. P. Jansson, *Deconvolution: With Applications in Spectroscopy* (Academic Press, 1995).
6. A. W. Howard, *et al.*, *Astrophys. J.* **696**, 75 (2009).
7. J. T. Wright, G. W. Marcy, R. P. Butler, S. S. Vogt, *Astrophys. J. Suppl. Ser.* **152**, 261 (2004).
8. M. A. C. Perryman, *et al.*, *Astron. Astrophys.* **323**, L49 (1997).
9. J. T. Wright, A. W. Howard, *Astrophys. J. Suppl. Ser.* **182**, 205 (2009).
10. J. A. Valenti, D. A. Fischer, *Astrophys. J. Suppl. Ser.* **159**, 141 (2005).
11. G. Takeda, *et al.*, *Astrophys. J. Suppl. Ser.* **168**, 297 (2007).
12. W. J. Borucki, *et al.*, *Science* **327**, 977 (2010).
13. W. J. Borucki, *et al.*, *ArXiv:1006.2799* (2010).
14. N. M. Batalha, *et al.*, *Astrophys. J.* **713**, L109 (2010).

15. R. P. Butler, *et al.*, *Astrophys. J.* **646**, 505 (2006).
16. D. A. Fischer, G. W. Marcy, R. P. Butler, G. Laughlin, S. S. Vogt, *Astrophys. J.* **564**, 1028 (2002).
17. M. Mayor, D. Queloz, *Nature* **378**, 355 (1995).
18. D. A. Fischer, *et al.*, *Astrophys. J.* **675**, 790 (2008).
19. S. S. Vogt, *et al.*, *Astrophys. J.* **708**, 1366 (2010).
20. G. W. Marcy, R. P. Butler, *Astrophys. J.* **464**, L147 (1996).
21. E. J. Rivera, *et al.*, *Astrophys. J.* **708**, 1492 (2010).
22. R. A. Wittenmyer, M. Endl, W. D. Cochran, H. F. Levison, G. W. Henry, *Astrophys. J.* **182**, L97 (2009).
23. C. Lovis, *et al.*, *Nature* **441**, 305 (2006).
24. D. Fischer, *et al.*, *Astrophys. J.* **703**, 1545 (2009).
25. HD 90156 b was announced by a press release (<http://www.eso.org/public/news/eso0939/>) in 2009 Oct. based on observations with HARPS. Our Keck/HIRES observations of this star confirm the properties listed in Table S2.
26. R. A. Wittenmyer, M. Endl, W. D. Cochran, H. F. Levison, G. W. Henry, *Astrophys. J. Suppl. Ser.* **182**, 97 (2009).
27. J. T. Wright, *et al.*, *Astrophys. J.* **683**, L63 (2008).
28. A. W. Howard, *et al.*, *Astrophysical Journal (submitted)*, *arXiv:1003.3444* (2010).
29. J. T. Wright, *et al.*, *Astrophys. J.* **693**, 1084 (2009).

30. S. S. Vogt, *et al.*, *Astrophys. J.* **632**, 638 (2005).